

Cosmologie

L. Vacher -

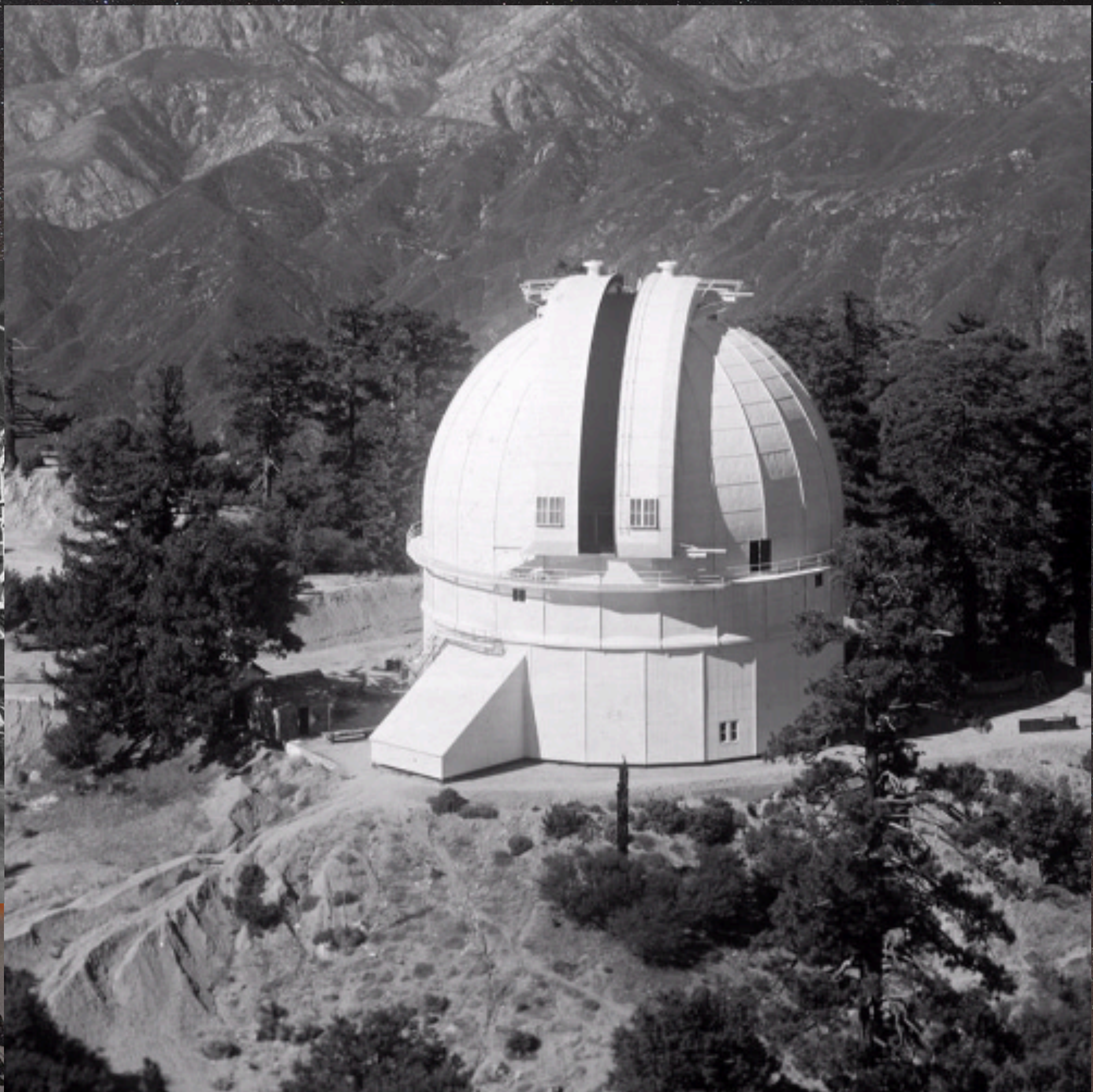
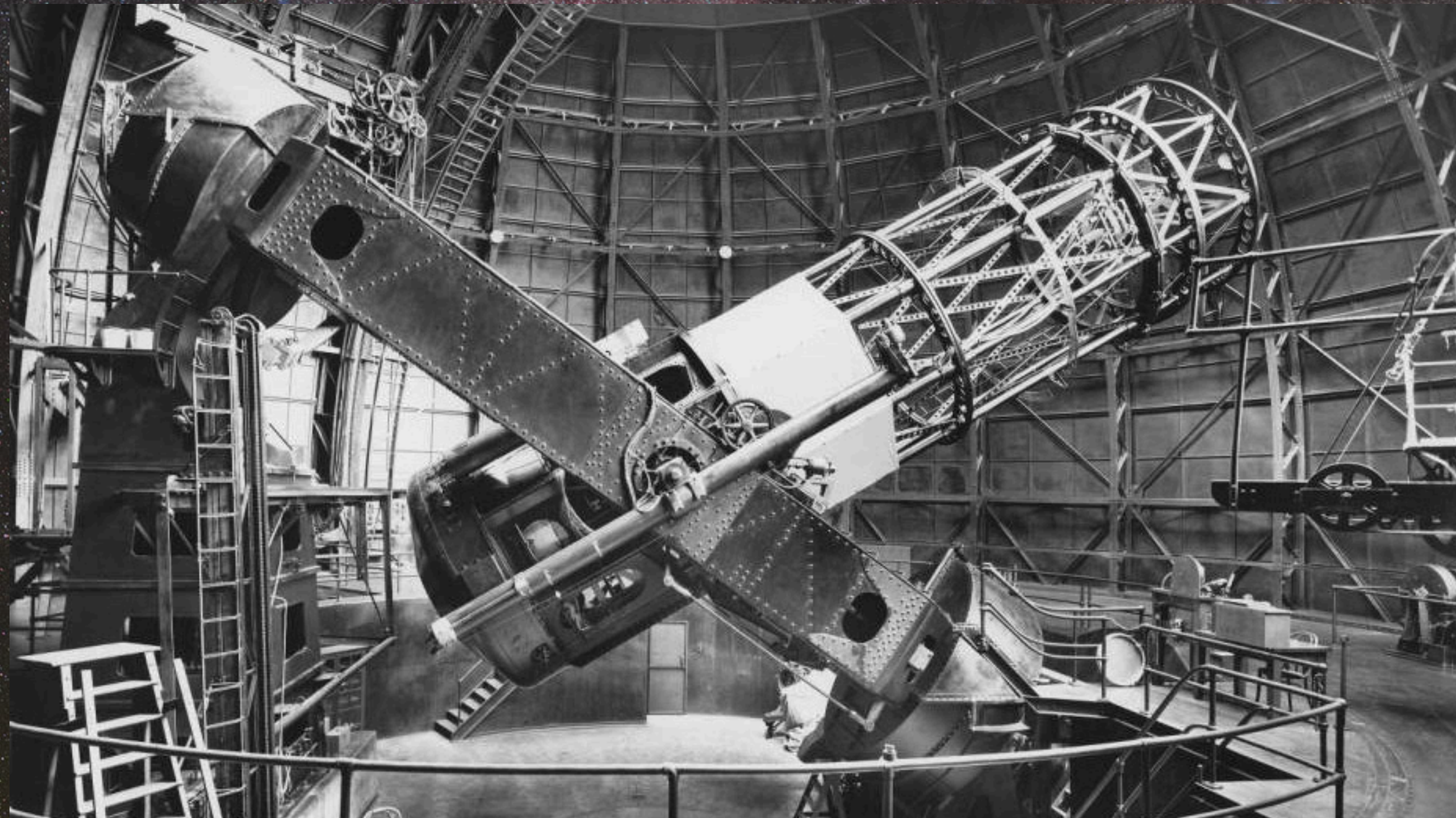


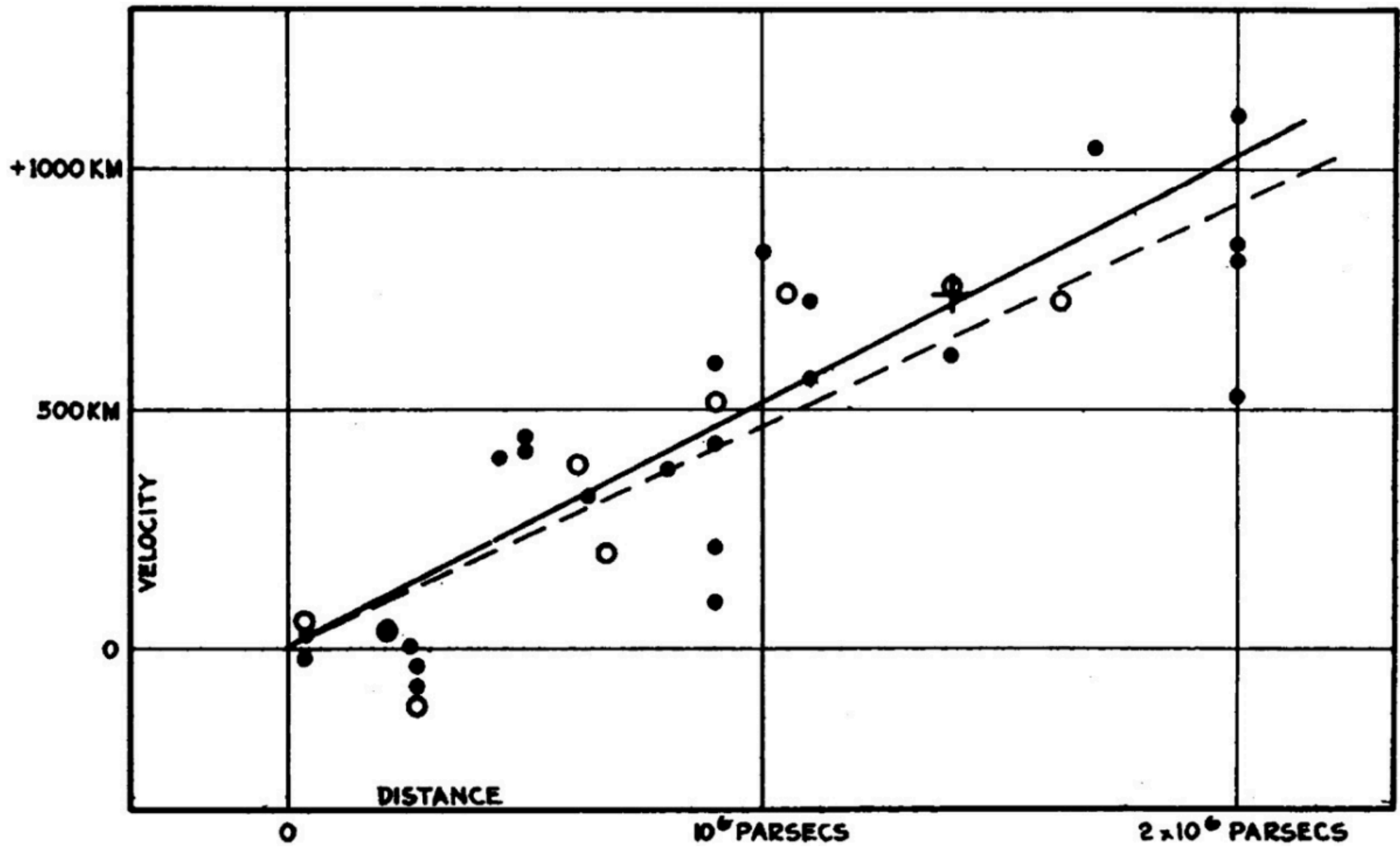
Edwin Hubble



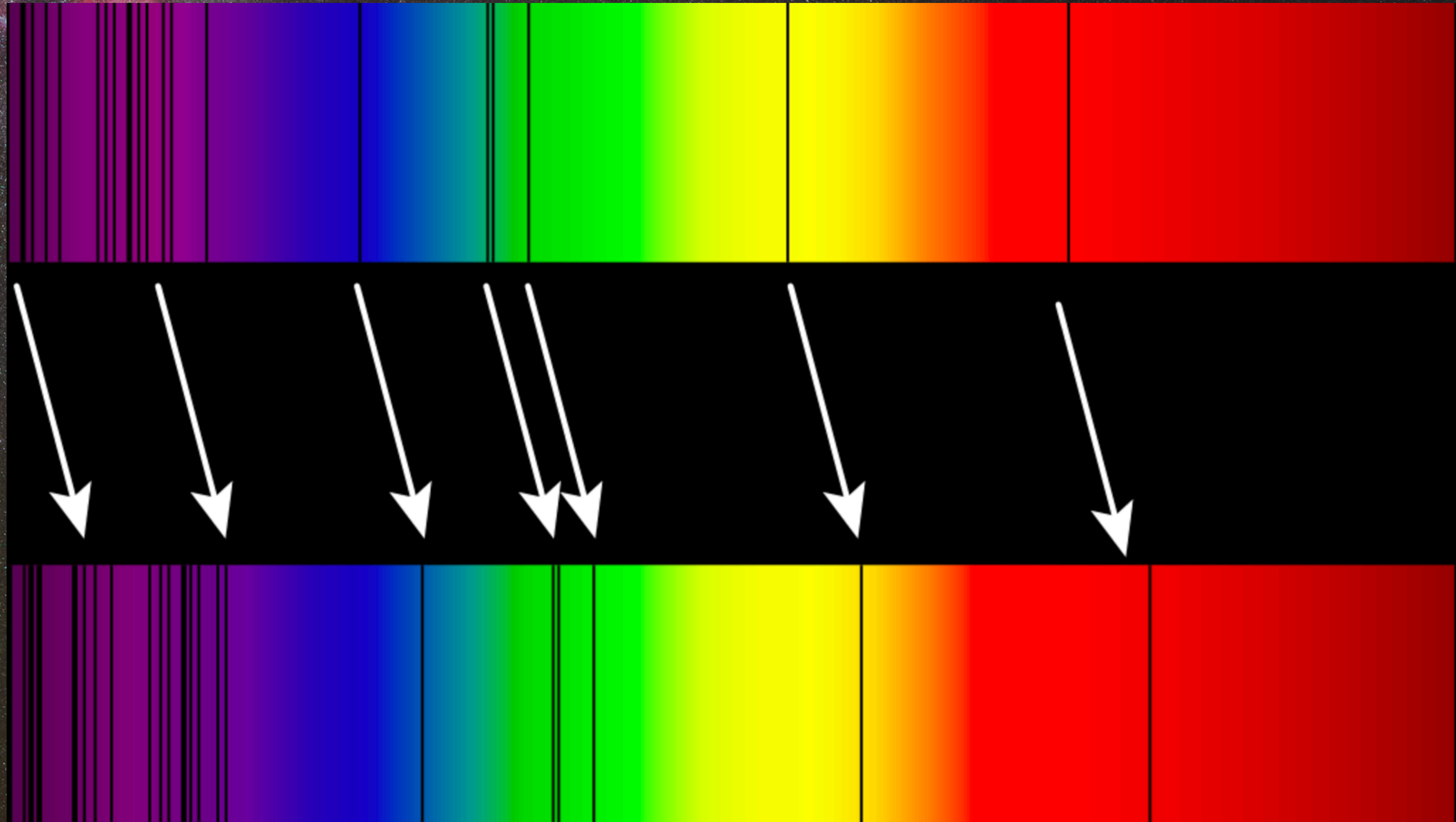
Milton Humason





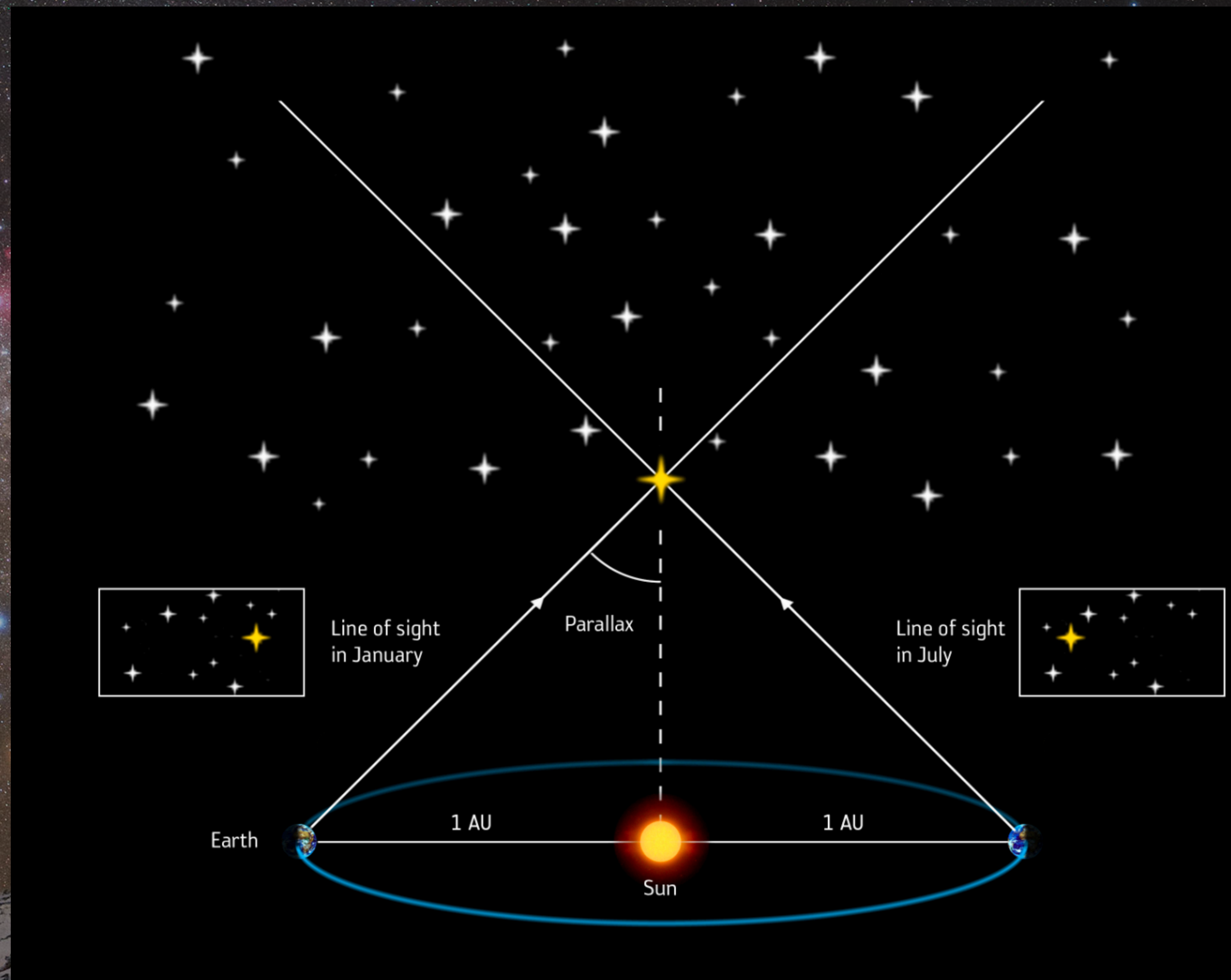


Vitesse?



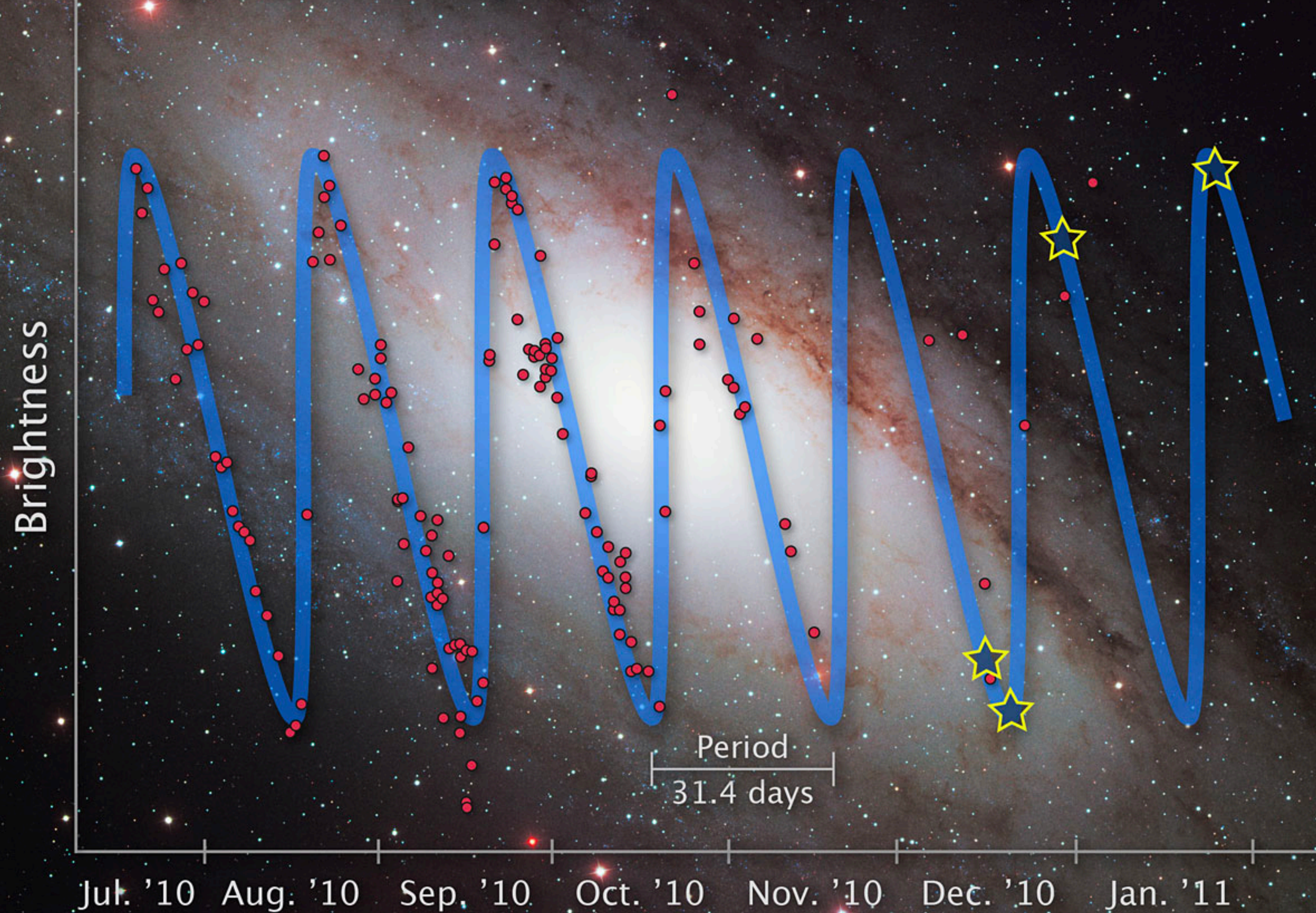
Distance?

Parallaxes ?

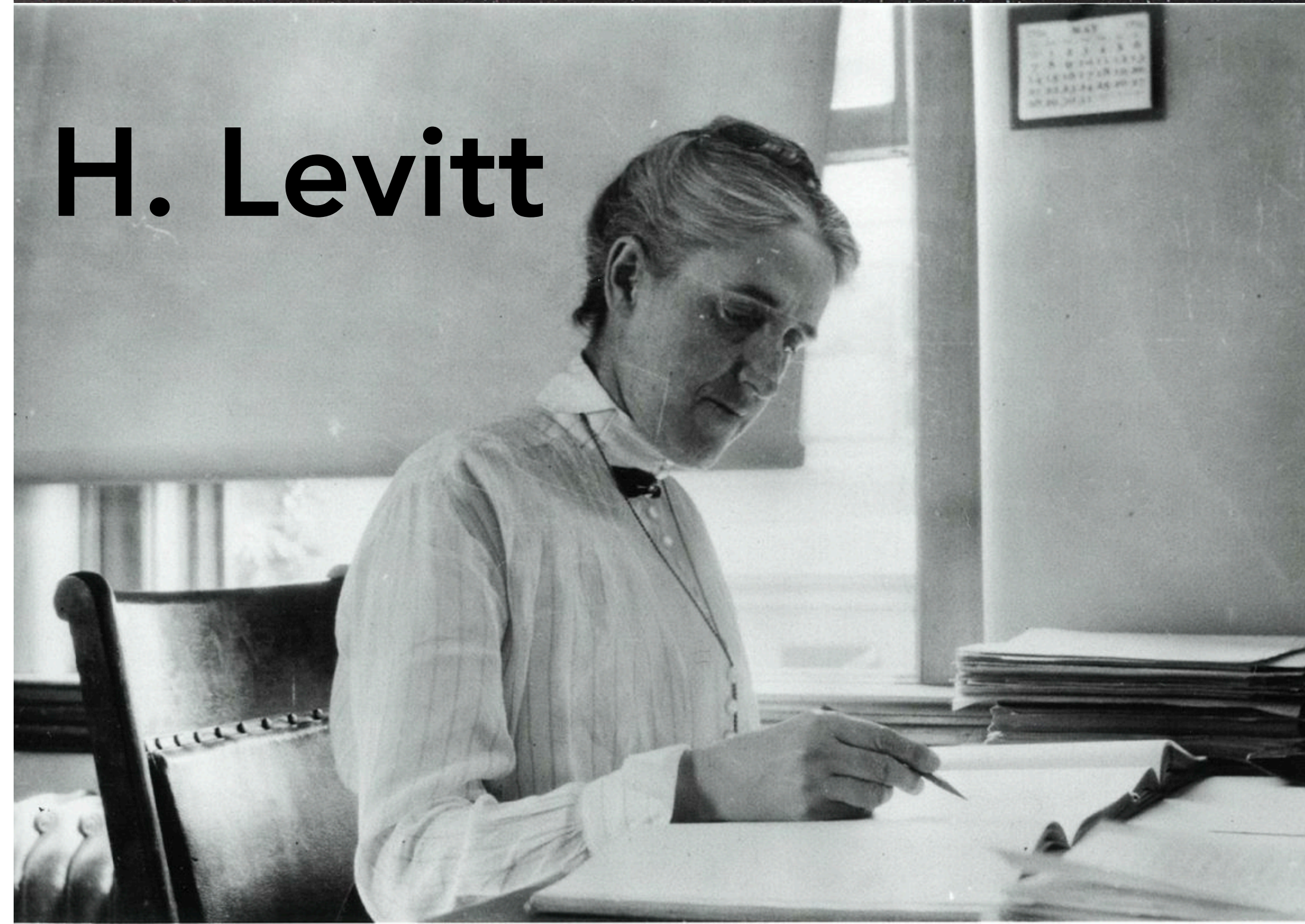
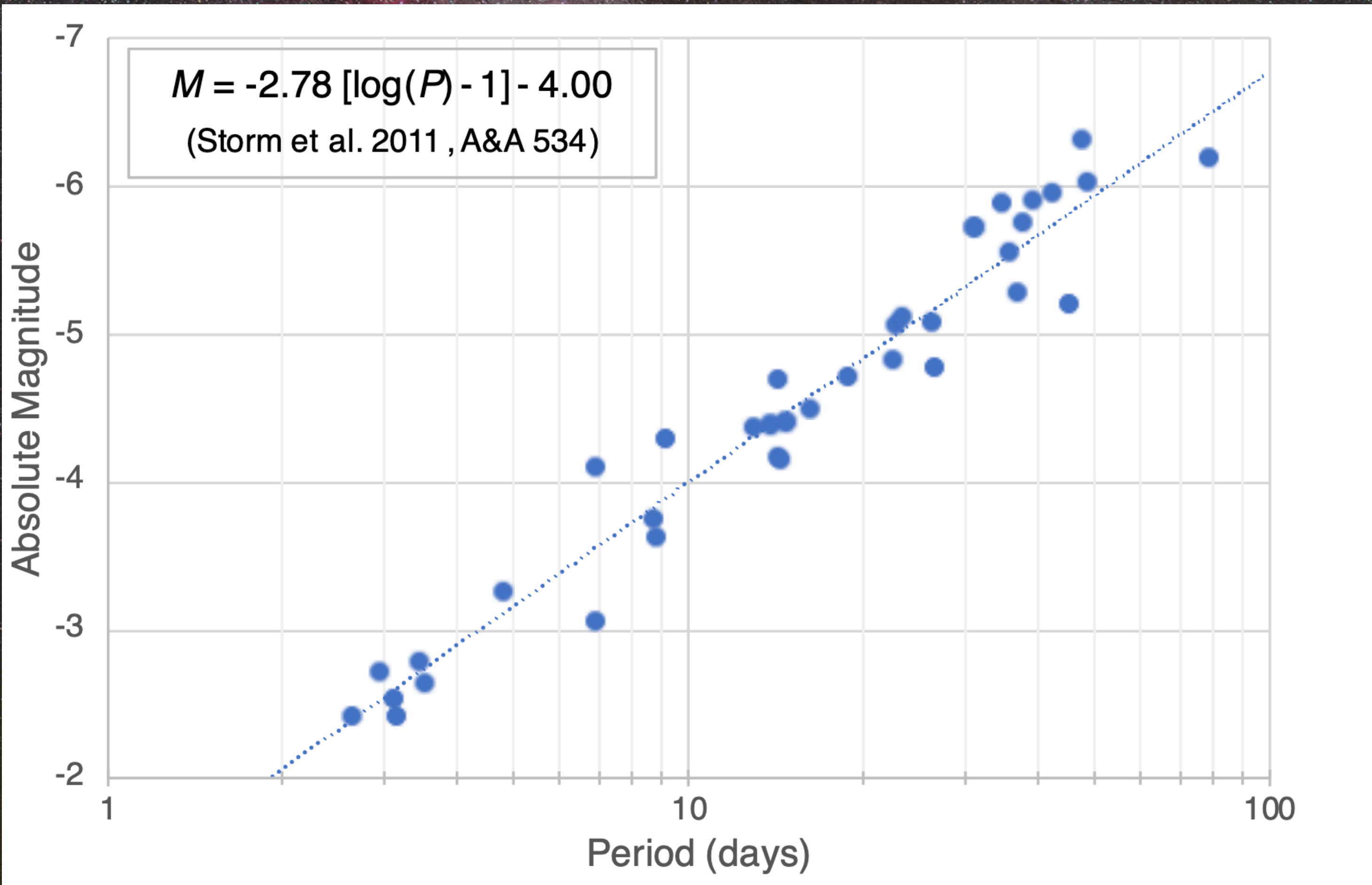


Distance? Céphéides

Light curve of Cepheid variable star V1 in galaxy M31



Distance? Céphéïdes



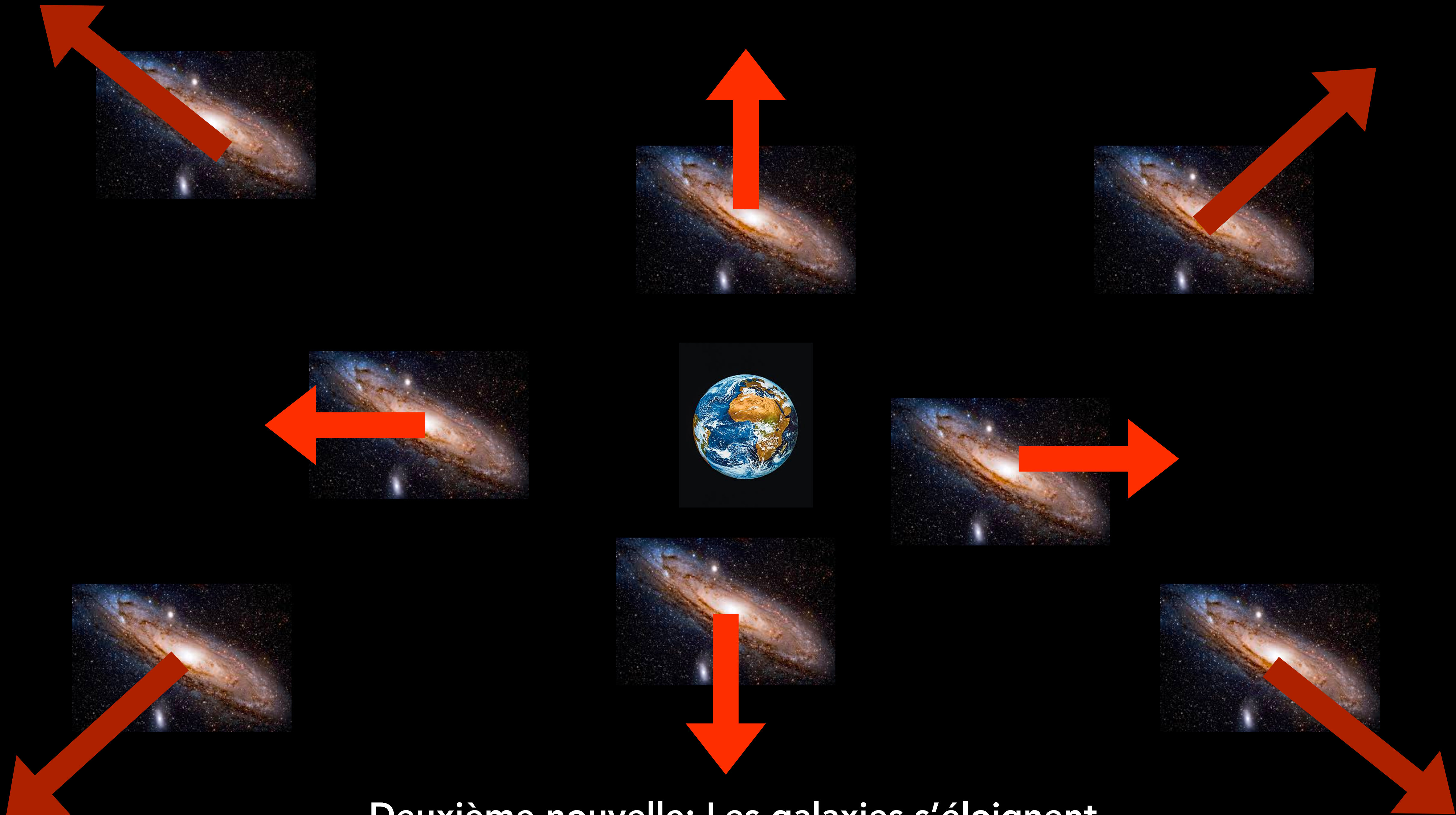
H. Levitt



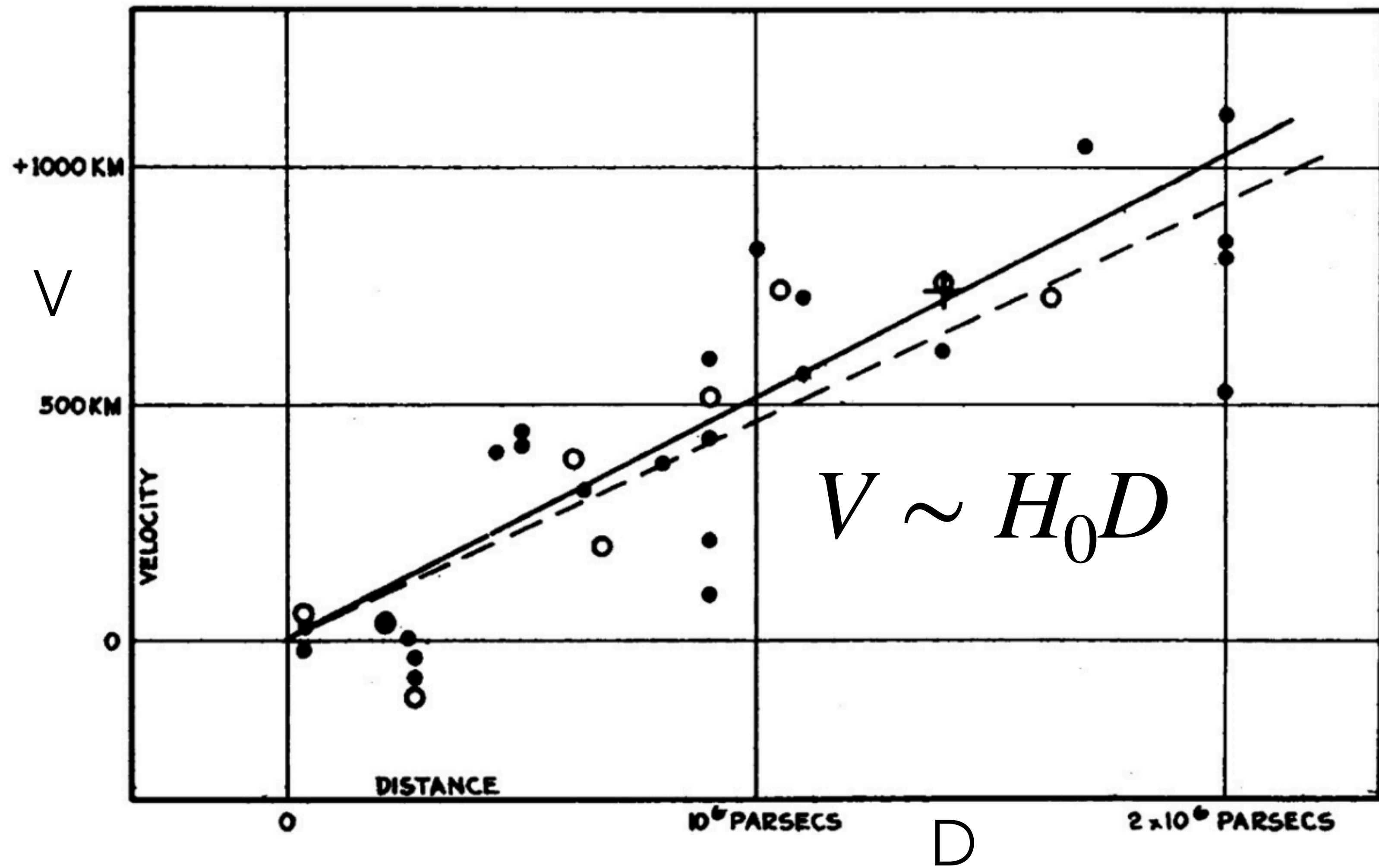
Première nouvelle: « nébuleuses » vs « galaxies »



Première nouvelle: « nébuleuses » vs « galaxies » (The Hubble extremely deep field)



Deuxième nouvelle: Les galaxies s'éloignent



Deuxième nouvelle: Les galaxies s'éloignent

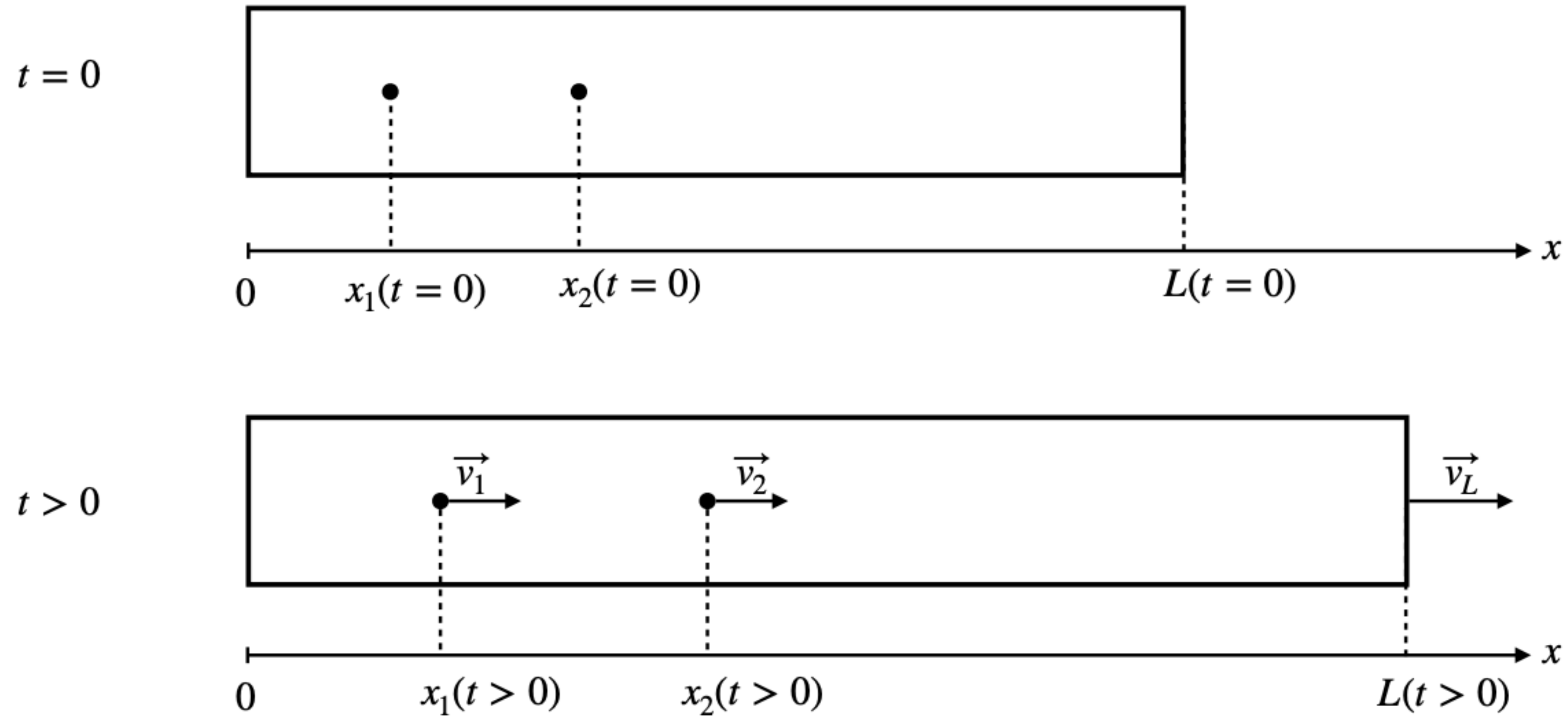
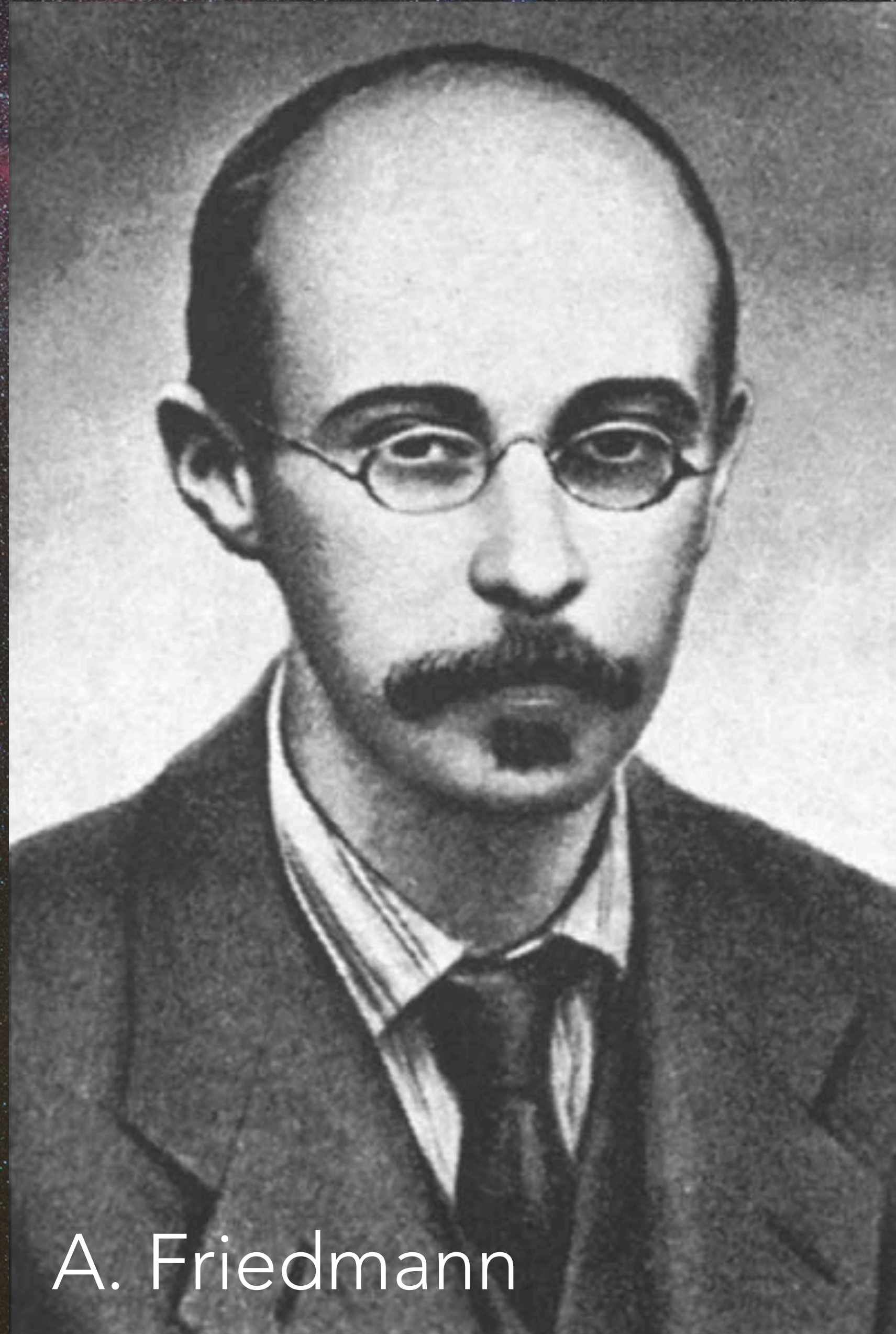


FIGURE 1 – Représentations schématiques de l'élastique en expansion à $t = 0$ et à $t > 0$



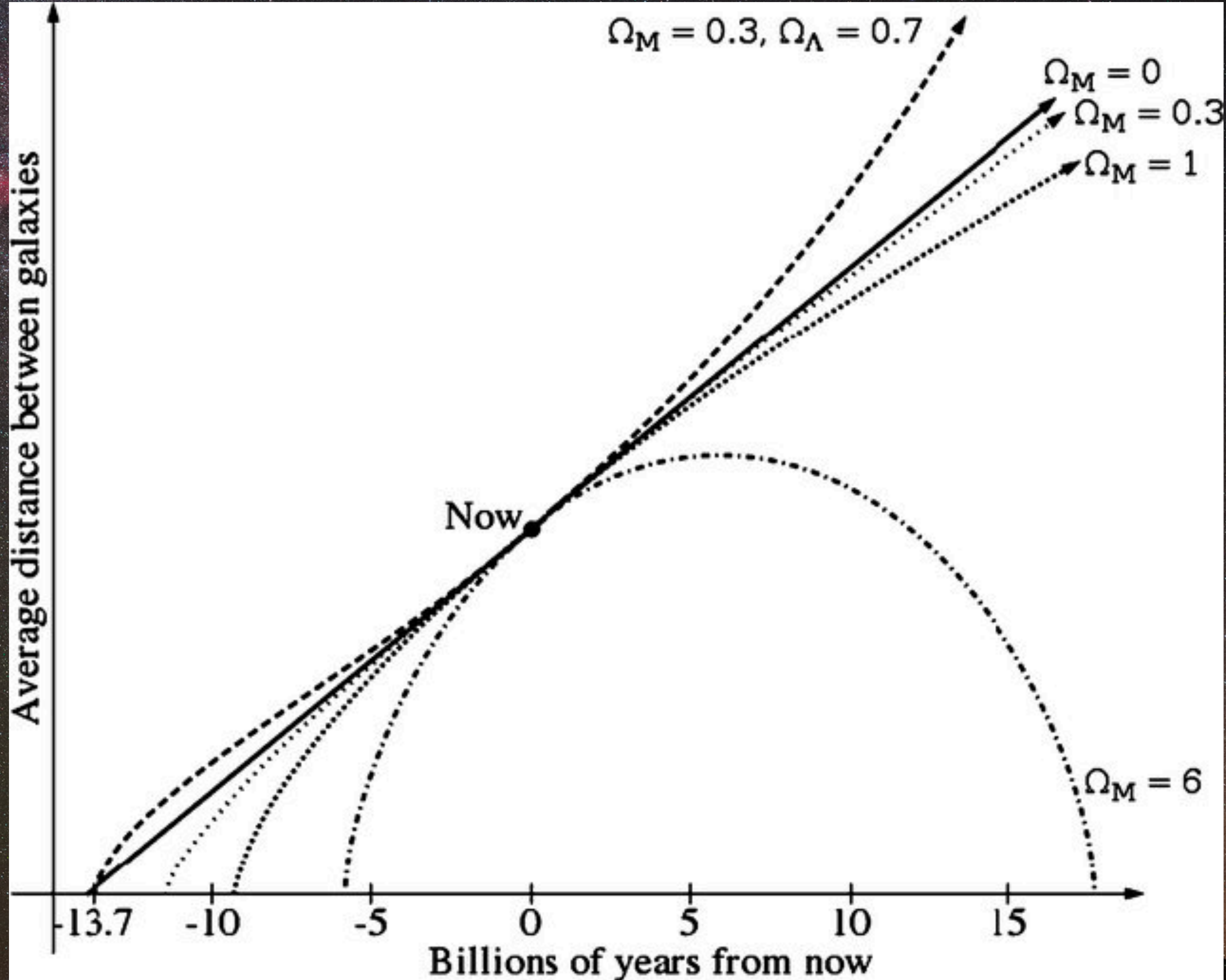
A. Friedmann



G. Lemaitre

A visualization of the cosmic web, showing a dense network of purple and orange filaments and nodes against a dark background. The central text 'Principe cosmologique' is overlaid in white. The image is framed by a dark, starry sky on the left and right sides.

Principe cosmologique



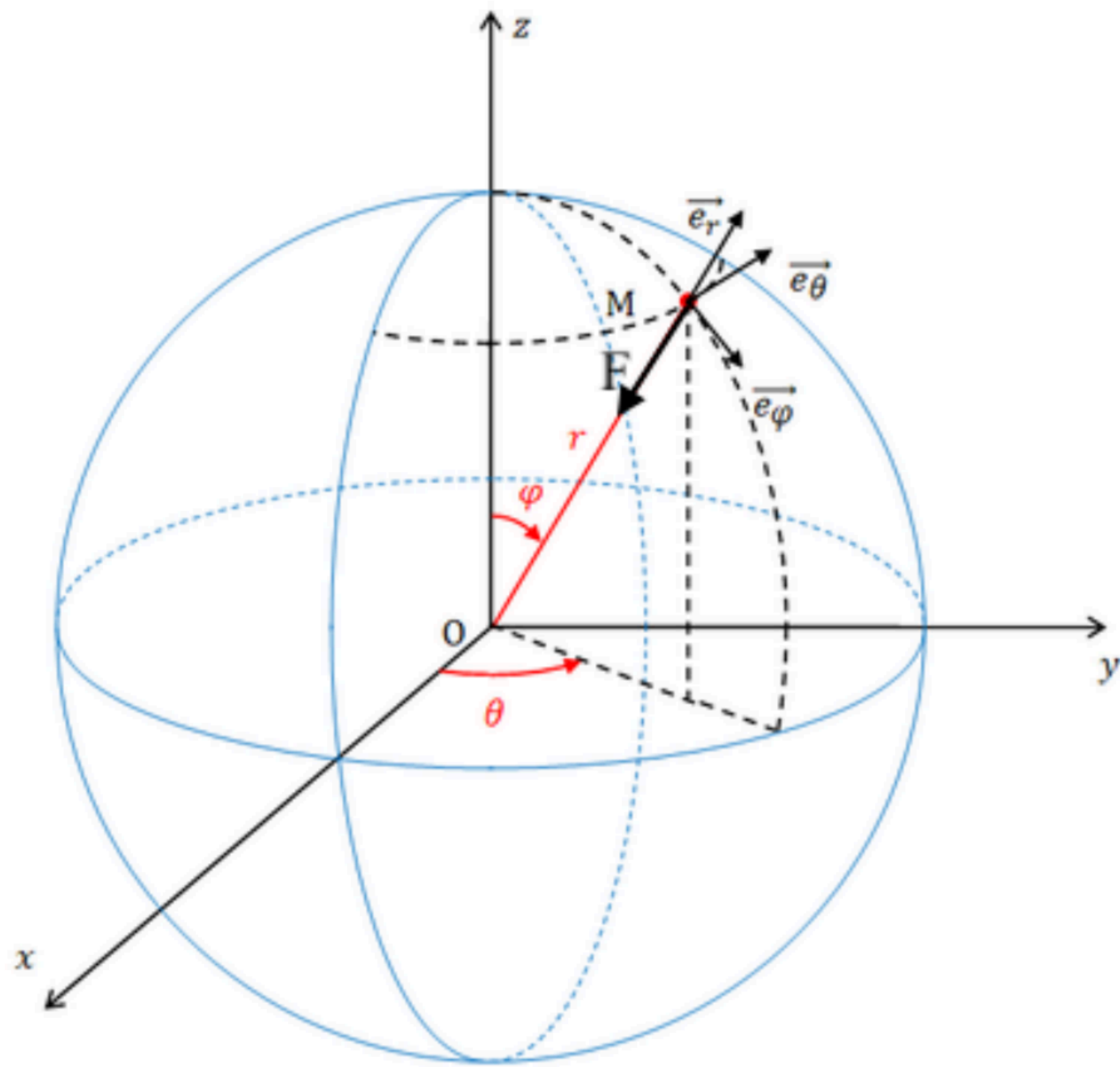
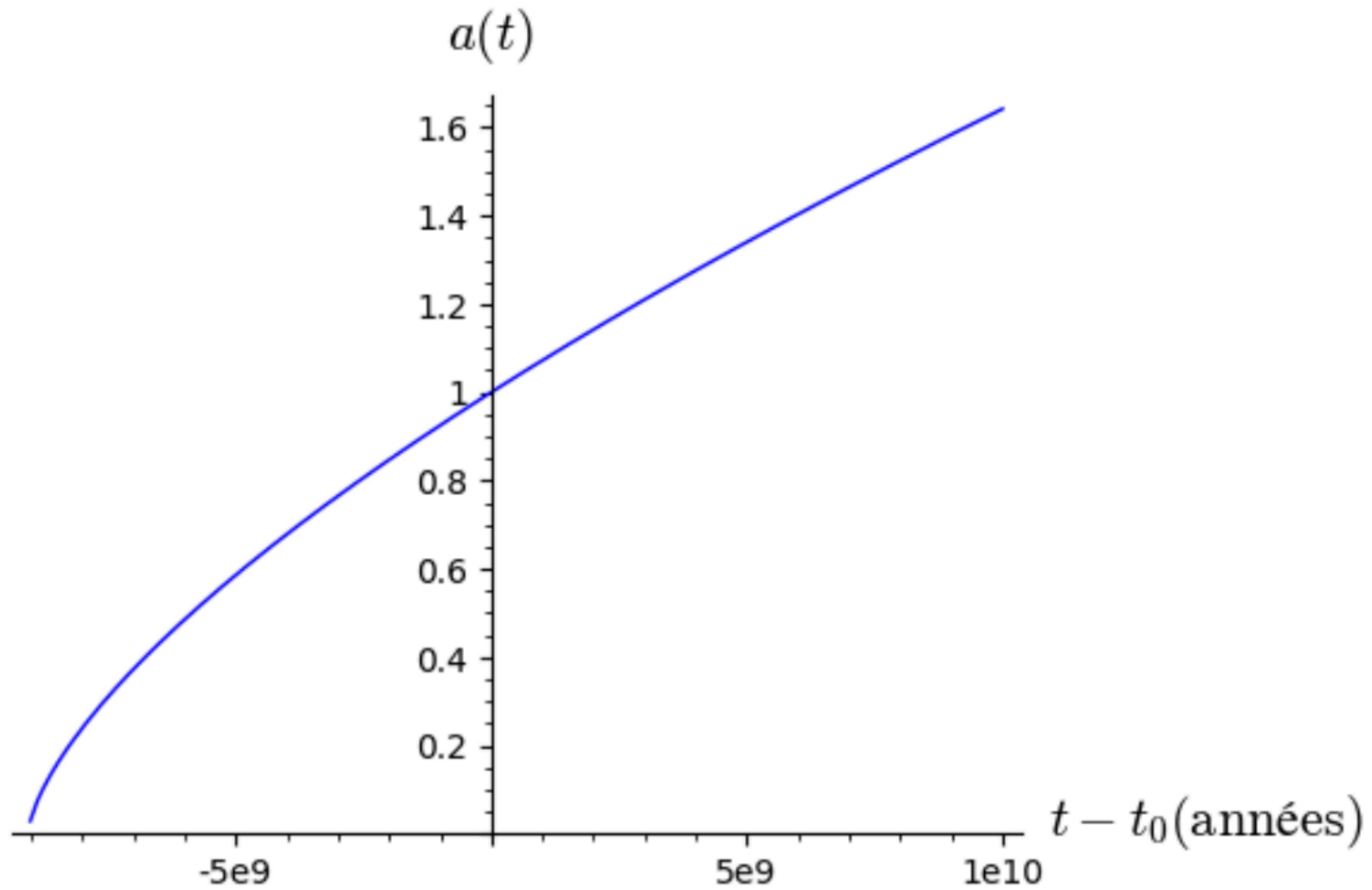
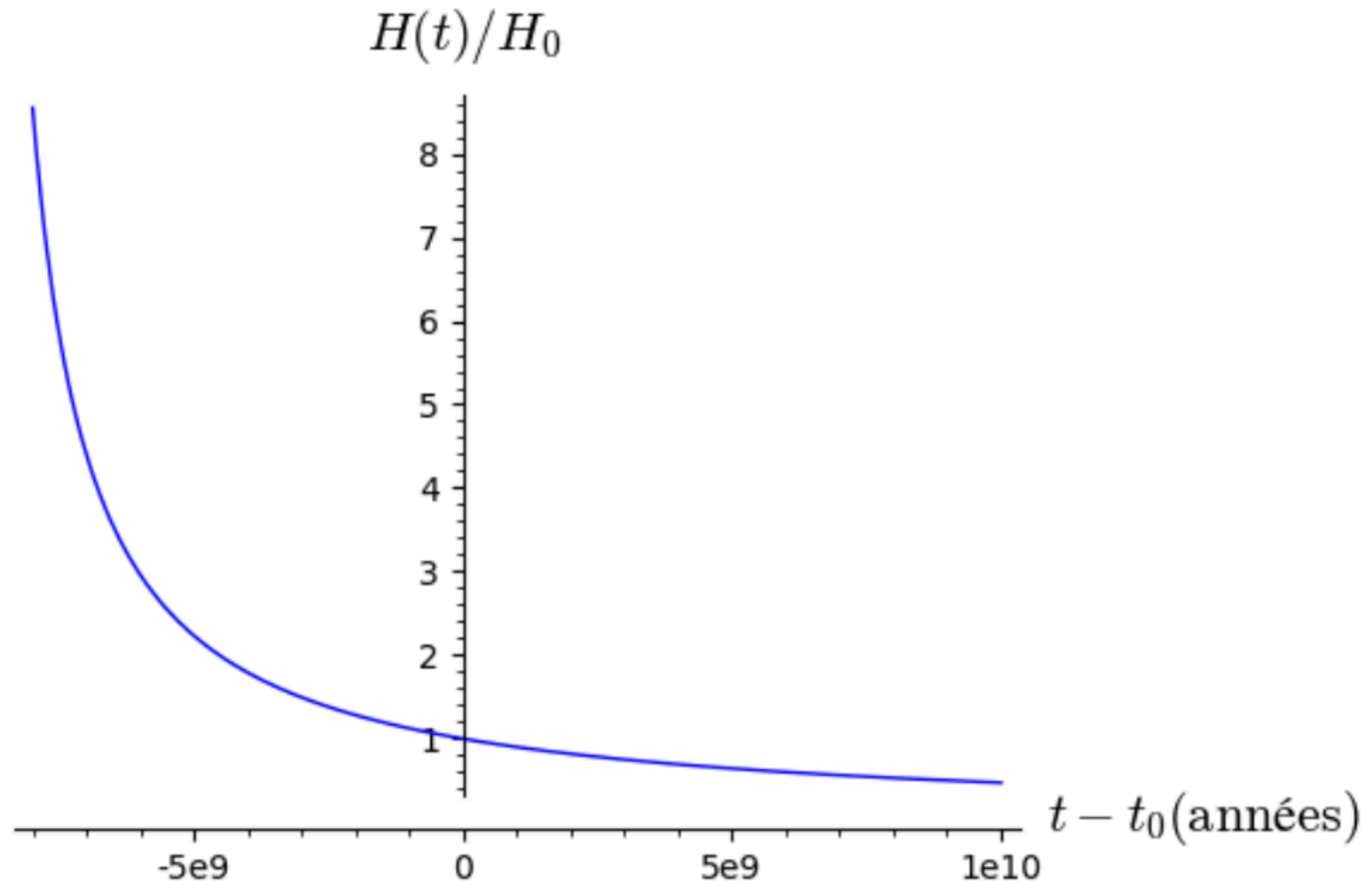


FIGURE 3 –





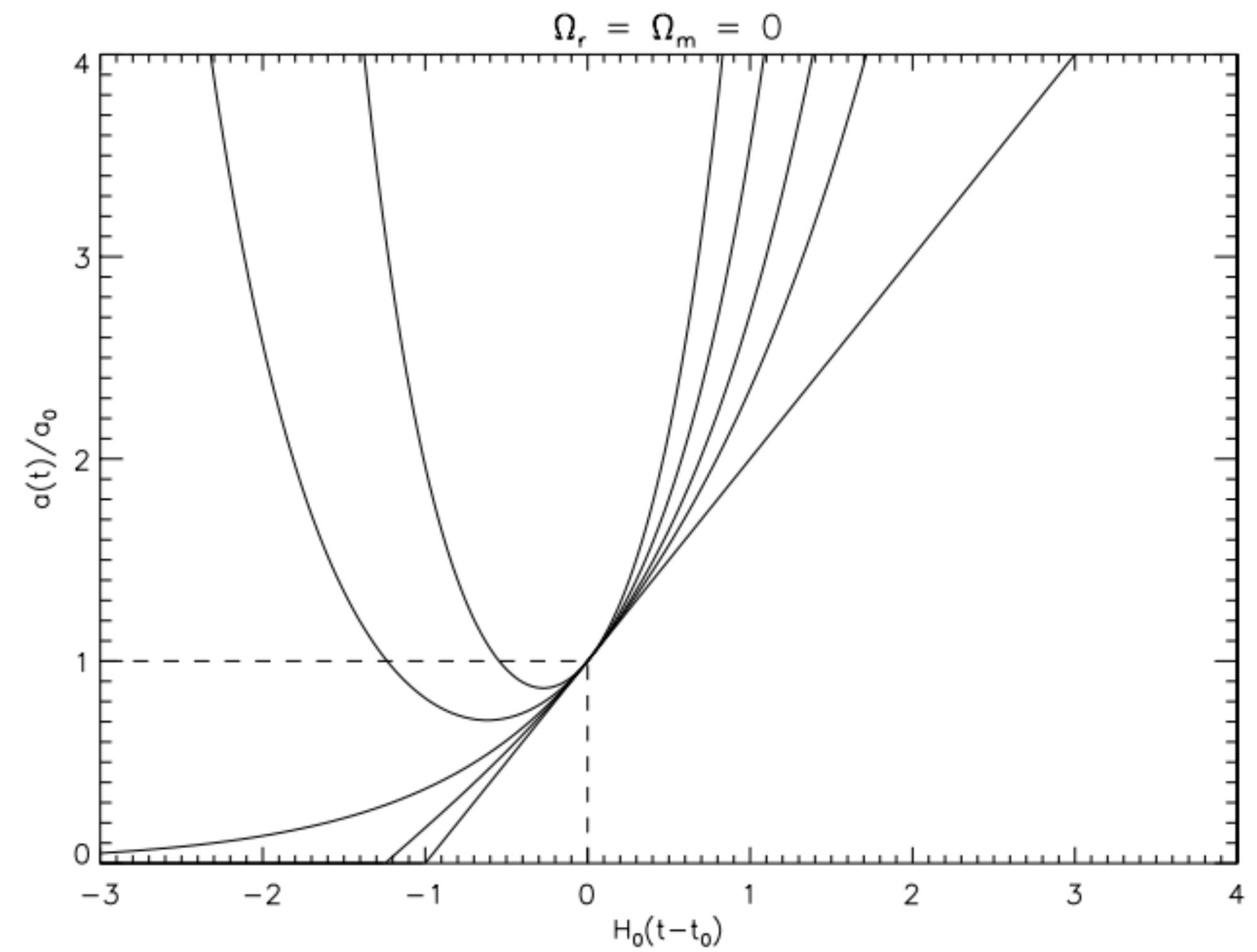
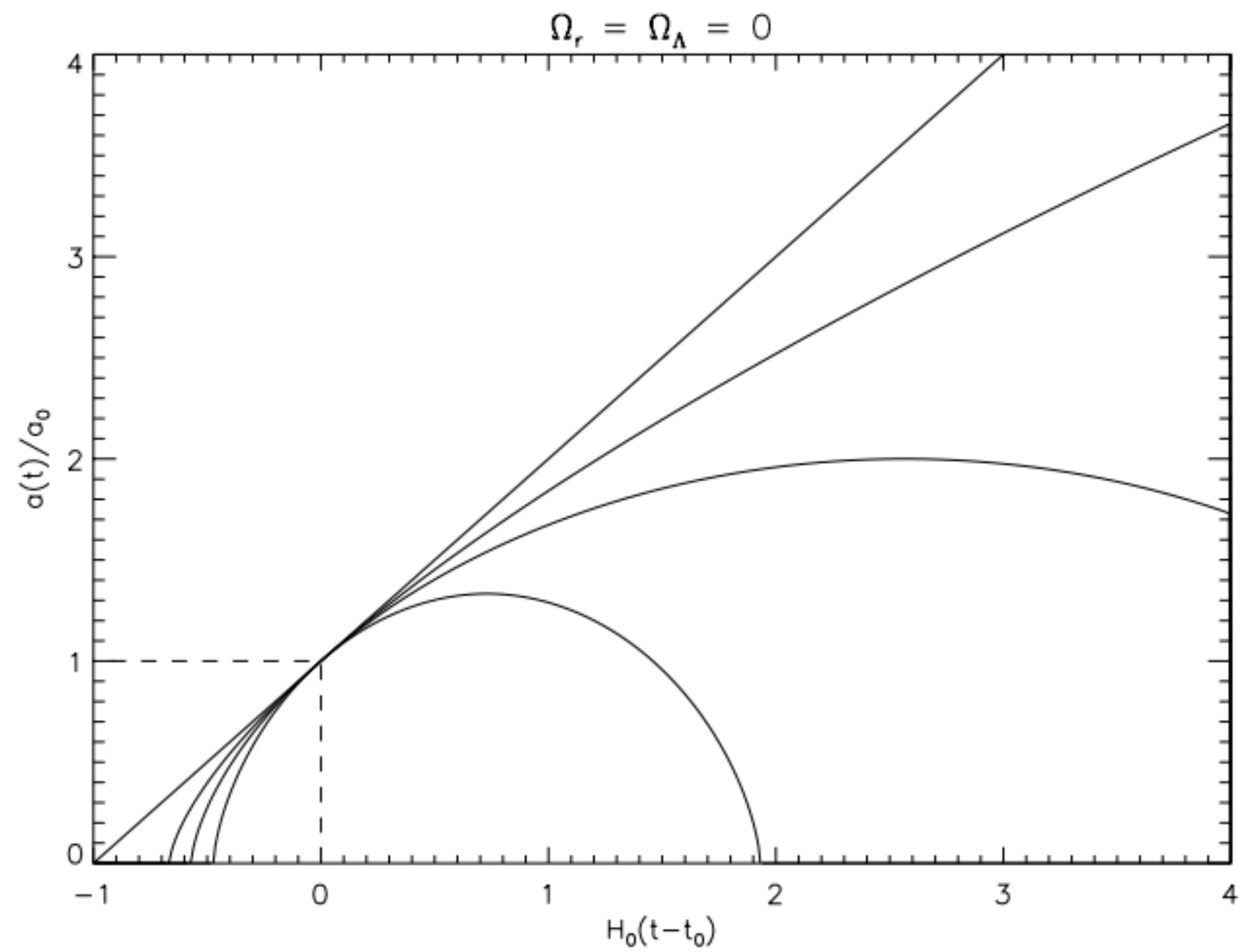
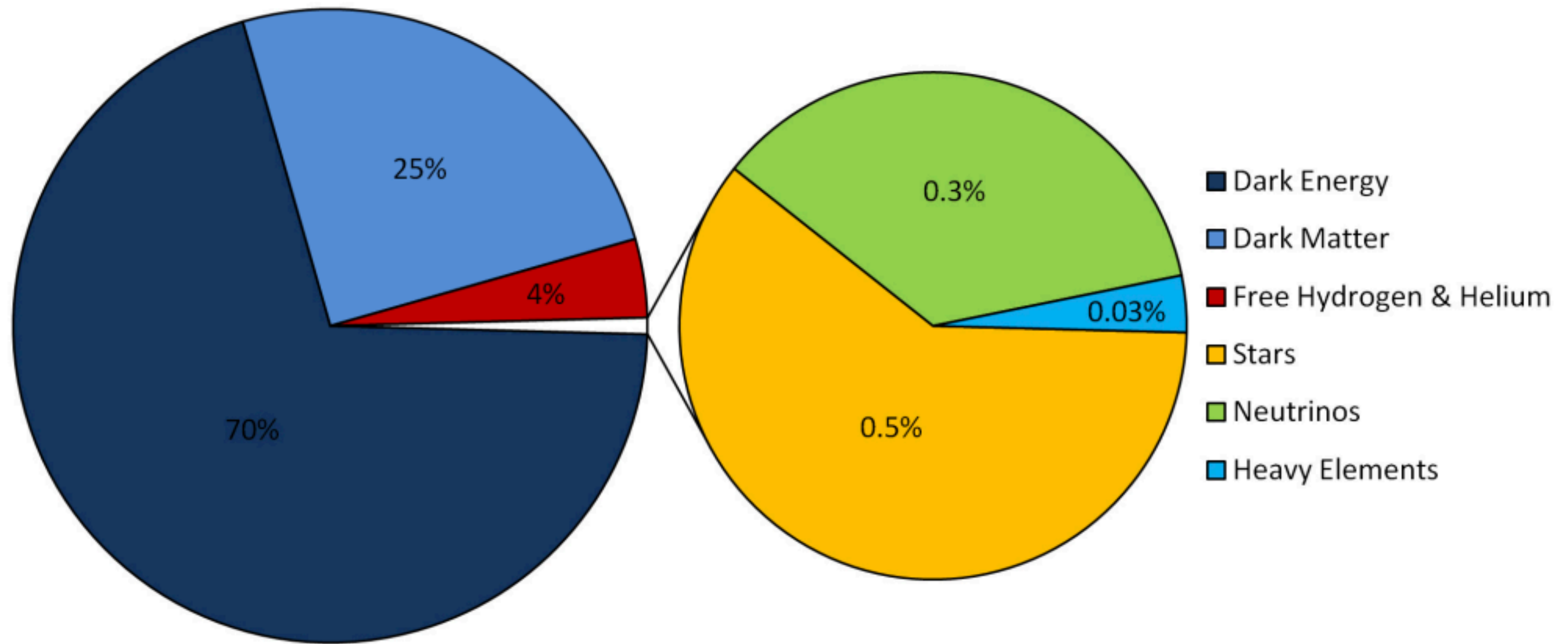
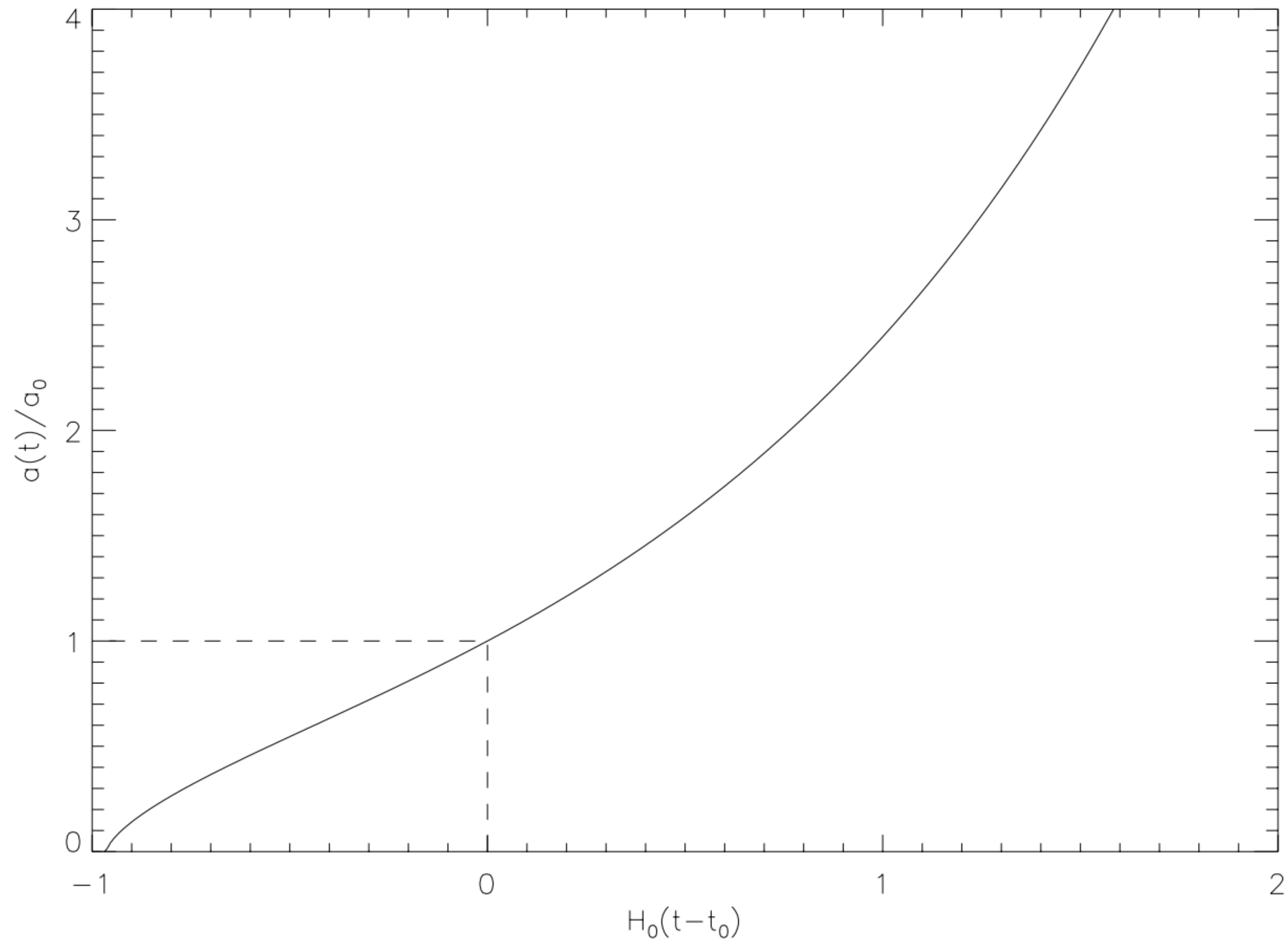
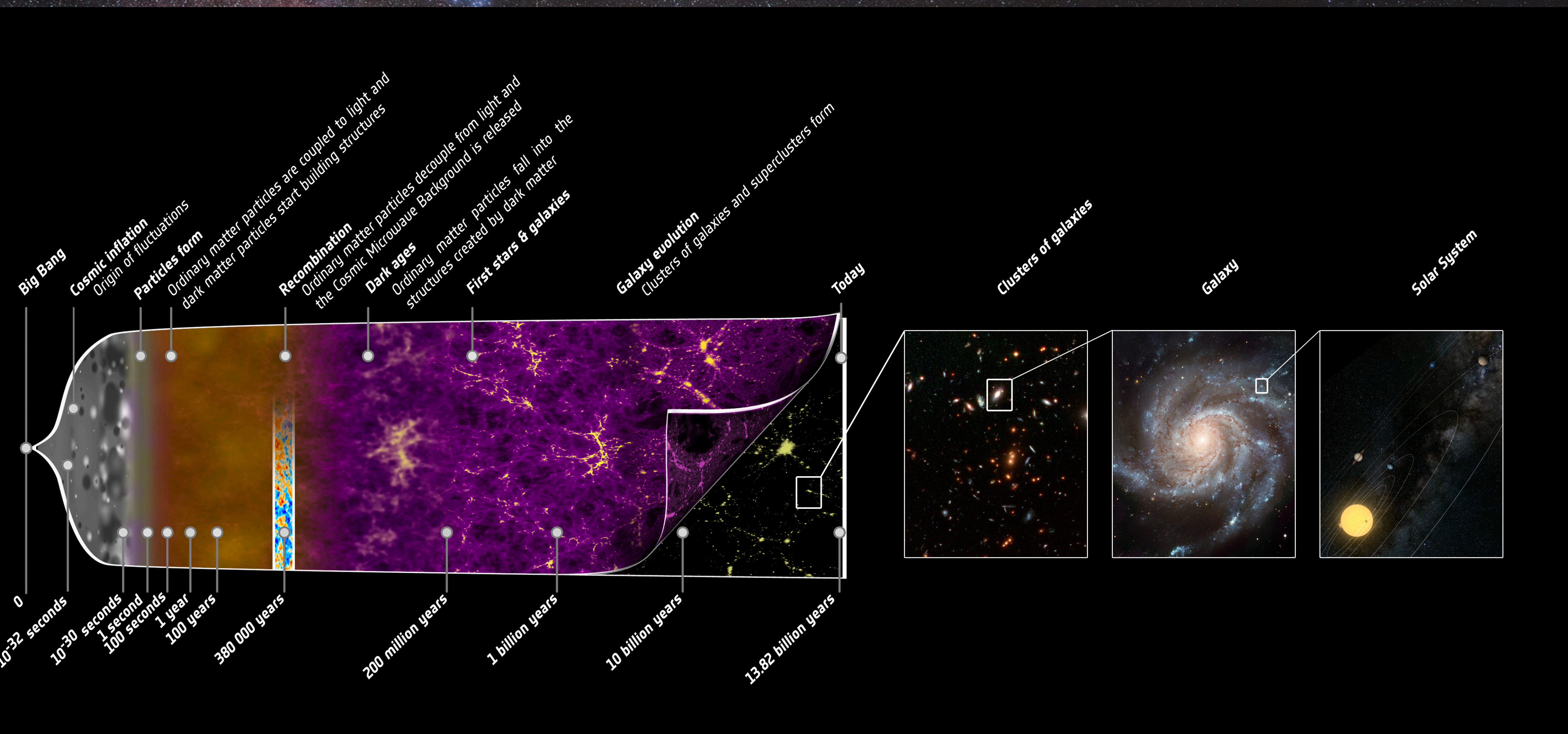
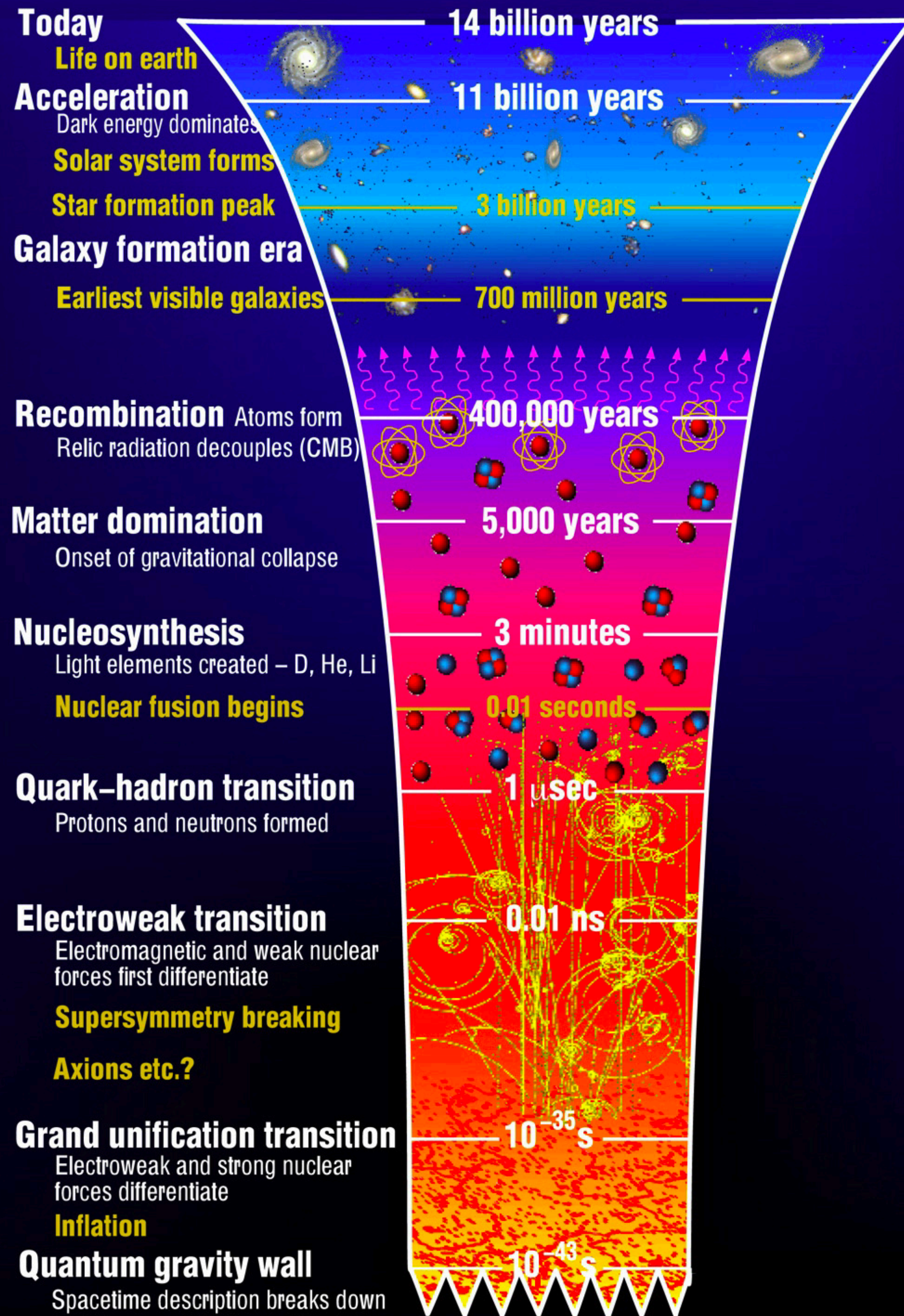


FIGURE 2 – Evolution de l'univers, représentée par son facteur d'échelle a en fonction du temps par rapport à l'époque actuelle. *A gauche* : Univers de matière ($\Omega_r = \Omega_\Lambda = 0$), pour $\Omega_m = 0, 1, 2, 4$. *A droite* : Univers de De Sitter ($\Omega_r = \Omega_m = 0$) pour $\Omega_\Lambda = 0, 0.5, 1, 2, 4$.









flat Λ CDM

